

# Féeton Dark Matter and probing high-energy physics with low-energy neutrino astronomy

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Publications:

- 1 Lin et al. [2022](#), PRD, arXiv: 2205.08171
- 2 Lin and Yanagida [2022](#), PRDL, arXiv: 2202.04496
- 3 Cheng et al. [2023](#), arXiv: 2310.05420

# What is féeton?

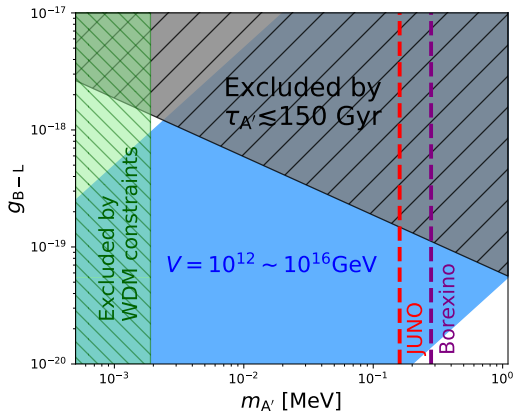
Féeton DM:  $B-L$  gauge boson with a tiny coupling  $10^{-19}$  and a sub-MeV mass.

	Electromagnetism:	B-L gauge (hypothesis):
Charge	Positive – Negative	Baryon number – Lepton number
Force mediator	Electromagnetic field	B-L gauge boson
Range	Long, massless boson	Short, massive gauge field
Why needed?	Observed	Motivated by neutrino mass, cosmic baryon asymmetry via <u>leptogenesis</u>
		A new DM candidate

The model is characterized by the gauge coupling  $g_{B-L}$  and mass  $m_f$ .

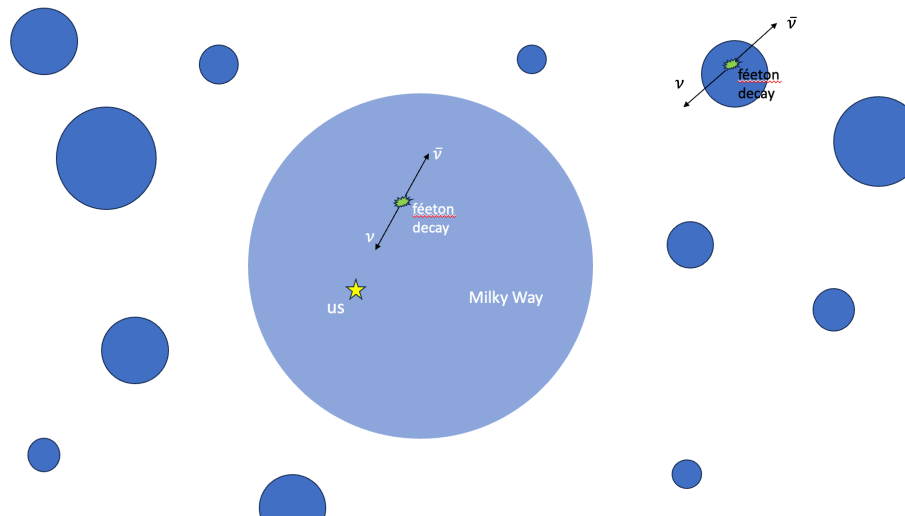
# The féeton miracle

- 1 Needs to be sub-MeV, otherwise strong constraints;
- 2 Lifetime needs to be sufficiently long (gray);
- 3 Mass above keV (green, but can be modified);
- 4 The Seesaw mechanism consistency (blue).

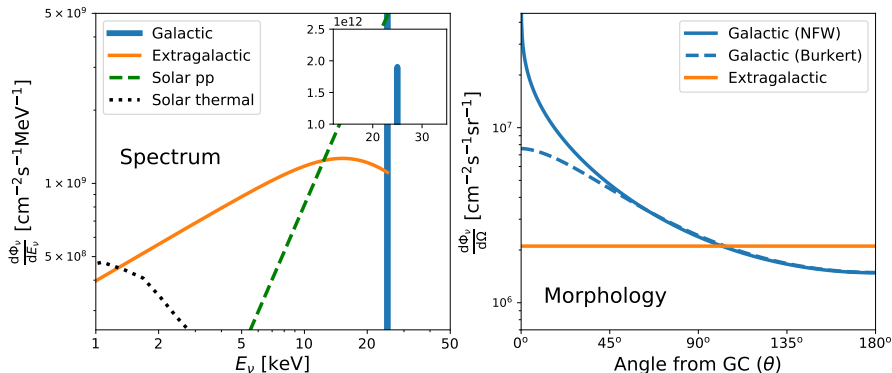


Miracle: if féeton is the dominant DM, the consistency with the Seesaw mechanism automatically makes it to be sub-MeV!

# Féeton decays into neutrinos



# Neutrino fluxes from Féeton decays



**Left:** Spectrum, taking maximal  $g_B$   $L$ , comparable to solar components at corresponding energy;

**Right:** Morphology, rather smooth but concentrated towards the Galactic center.

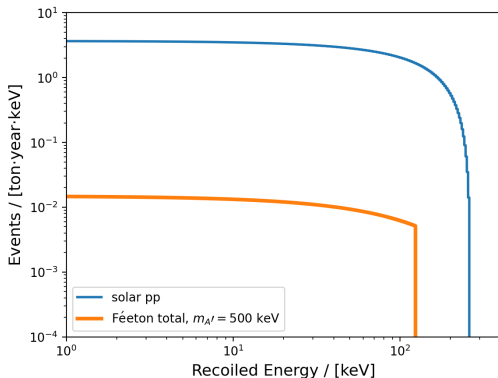
# What can we learn if we detect the féeton neutrinos

$$\frac{d\Phi_\nu^{\text{Gal}}}{d\Omega} = 7.9 \times 10^5 \tilde{D}_N(\cos\theta) \left(\frac{g_{B-L}}{10^{-19}}\right)^2 [\text{cm}^{-2} \text{s}^{-1} \text{sr}^{-1}]$$

- The energy of the neutrino  $\sim$  féeton mass;
- Total flux  $\sim$   $B-L$  gauge coupling constant  $g_{B-L}$ ;
- Then from  $V_{B-L} = m_f/2g_{B-L}$   $\sim$   $B-L$  symmetry breaking scale. This provides a unique way to probe the high-energy  $B-L$  physics.

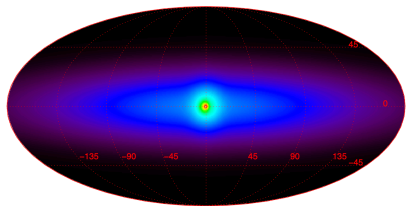
# Solar neutrino-type experiments

- Solar pp neutrinos dominate;
- Direction determination will be helpful;
- JUNO and future Broxino can survey  $\lesssim 100$  keV.



# What if $m_f > 2m_e$ ? – féeton in addition decays into electron-positron pairs

## Galactic 511-keV line (positron annihilation) excess



From INTEGRAL (Siegert et al. 2016)

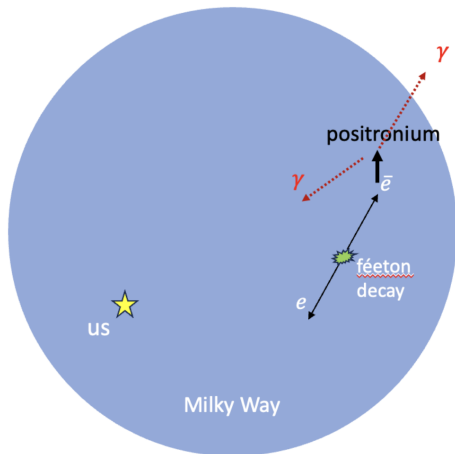
Total	2.74	0.25	$10^{-3} \text{ cm}^{-1} \text{ s}^{-1}$
Bulge	0.96	0.07	$10^{-3} \text{ cm}^{-1} \text{ s}^{-1}$
Disk	1.66	0.35	$10^{-3} \text{ cm}^{-1} \text{ s}^{-1}$
$f_{Ps}$	1		
$K_{ini}$	3 MeV		

Adopted from (Siegert et al. 2016).

- First discovered in 1970's,  $> 50$  years;  
(Johnson, Harnden, and Haymes 1972; Haymes et al. 1975; Leventhal, MacCallum, and Stang 1978)
- Mainly positronium (intermediate form);  
(Harris et al. 1998; Siegert et al. 2016; Jean et al. 2006)
- Sharp towards the Galactic center;  
(Ascasibar et al. 2006)
- Positron initial energy  $\sim 3$  MeV.  
(Beacom and Yüksel 2006; Sizun, Casse, and Schanne 2006)



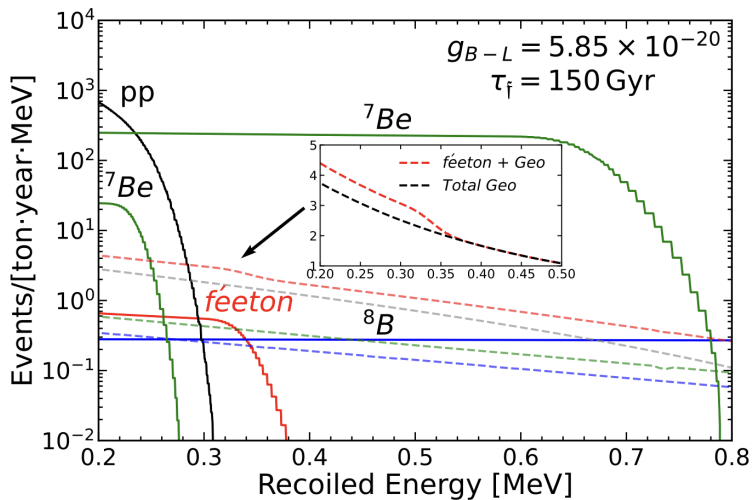
# Féeton decays and Galactic $\gamma$ -ray excess



Two possible processes:

- $m_f > 2m_e + 13.6 \text{ keV}$ :
  1. Form positronium via charge exchanges hydrogen atoms;
  2. féeton cannot be dominant DM; Lin and Yanagida [2022](#), PRDL
- $2m_e < m_f < 2m_e + 13.6 \text{ keV}$ :
  1. Form positronium with free electrons in ionized environment;
  2. féeton can be dominant DM; Cheng et al. [2023](#), arXiv: 2310.05420

# 511 keV positron annihilation anomaly



# Summary

- $B-L$  gauge boson with a tiny coupling  $g_{B-L} \sim 10^{-19}$  and a sub-MeV mass;
- Consistent in cosmology, seesaw mechanism, and leptogenesis;
- A low-energy neutrino signal potentially detectable in the near future;
- *Fée*, a French word for fairy — a small being that carries an important message from heaven.






- Inflationary vector boson production (Graham, Mardon, and Rajendran 2016)

$$\Omega_X = \Omega_{\text{cdm}} \sqrt{\frac{m_f}{6 \cdot 10^9 \text{ keV}}} \left( \frac{H_{\text{inf}}}{10^{14} \text{ GeV}} \right)^2$$






Then,  $H_{\text{inf}} \approx 1.6 \cdot 10^{11} \text{ GeV}$  for  $m_f = 1 \text{ MeV}$  and  $\Omega_f = \Omega_X = \Omega_{\text{cdm}}$ .

- Recall  $H_{\text{inf}} \propto r^{1/2} \approx 2.6 \cdot 10^{12} \text{ GeV}$  for  $r = 4 \cdot 10^{-7}$ ;  
Falsified if PGWs should be detected in near future.




# References I

-  Ascasibar, Yago et al. (2006). “Constraints on dark matter and the shape of the Milky Way dark halo from the 511-keV line”. In: *Mon. Not. Roy. Astron. Soc.* 368, pp. 1695–1705. DOI: [10.1111/j.1365-2966.2006.10226.x](https://doi.org/10.1111/j.1365-2966.2006.10226.x). arXiv: [astro-ph/0507142](https://arxiv.org/abs/astro-ph/0507142).
-  Beacom, John F. and Hasan Yuksel (2006). “Stringent constraint on galactic positron production”. In: *Phys. Rev. Lett.* 97, p. 071102. DOI: [10.1103/PhysRevLett.97.071102](https://doi.org/10.1103/PhysRevLett.97.071102). arXiv: [astro-ph/0512411](https://arxiv.org/abs/astro-ph/0512411).
-  Cheng, Yu et al. (Oct. 2023). “Féeton ( $B-L$  Gauge Boson) Dark Matter for the 511-keV Gamma-Ray Excess and the Prediction of Low-energy Neutrino Flux”. In: arXiv: [2310.05420](https://arxiv.org/abs/2310.05420) [hep-ph].
-  Graham, Peter W., Jeremy Mardon, and Surjeet Rajendran (2016). “Vector Dark Matter from Inflationary Fluctuations”. In: *Phys. Rev. D* 93.10, p. 103520. DOI: [10.1103/PhysRevD.93.103520](https://doi.org/10.1103/PhysRevD.93.103520). arXiv: [1504.02102](https://arxiv.org/abs/1504.02102) [hep-ph].
-  Harris, M. J. et al. (1998). “Transient Gamma-Ray Spectrometer Measurements of the Positron Annihilation Spectrum from the Galactic Center”. In: *Astrophys. J. Lett.* 501, p. L55. DOI: [10.1086/311429](https://doi.org/10.1086/311429). arXiv: [astro-ph/9803247](https://arxiv.org/abs/astro-ph/9803247).

# References II

-  Haymes, R. C. et al. (Nov. 1975). “Detection of nuclear gamma rays from the galactic center region.”. In: *Astrophys. J.* 201, pp. 593–602. DOI: [10.1086/153925](https://doi.org/10.1086/153925).
-  Jean, Pierre et al. (2006). “Spectral analysis of the galactic  $e^+ e^-$  annihilation emission”. In: *Astron. Astrophys.* 445, pp. 579–589. DOI: [10.1051/0004-6361:20053765](https://doi.org/10.1051/0004-6361:20053765). arXiv: [astro-ph/0509298](https://arxiv.org/abs/astro-ph/0509298).
-  Johnson W. N., III, Jr. Harnden F. R., and R. C. Haymes (Feb. 1972). “The Spectrum of Low-Energy Gamma Radiation from the Galactic-Center Region.”. In: *Astrophys. J. Lett.* 172, p. L1. DOI: [10.1086/180878](https://doi.org/10.1086/180878).
-  Leventhal, M., C. J. MacCallum, and P. D. Stang (Oct. 1978). “Detection of 511 keV positron annihilation radiation from the galactic center direction.”. In: *Astrophys. J. Lett.* 225, pp. L11–L14. DOI: [10.1086/182782](https://doi.org/10.1086/182782).
-  Lin, Weikang et al. (2022). “Neutrino astronomy as a probe of physics beyond the Standard Model: Decay of sub-MeV B-L gauge boson dark matter”. In: *Phys. Rev. D* 106.7, p. 075011. DOI: [10.1103/PhysRevD.106.075011](https://doi.org/10.1103/PhysRevD.106.075011). arXiv: [2202.04496](https://arxiv.org/abs/2202.04496) [hep-ph].

# References III

-  Lin, Weikang and Tsutomu T. Yanagida (2022). “Confronting the Galactic 511 keV emission with B-L gauge boson dark matter”. In: *Phys. Rev. D* 106.7, p. 075012. DOI: [10.1103/PhysRevD.106.075012](https://doi.org/10.1103/PhysRevD.106.075012). arXiv: [2205.08171](https://arxiv.org/abs/2205.08171) [hep-ph].
-  Siegert, Thomas et al. (2016). “Gamma-ray spectroscopy of Positron Annihilation in the Milky Way”. In: *Astron. Astrophys.* 586, A84. DOI: [10.1051/0004-6361/201527510](https://doi.org/10.1051/0004-6361/201527510). arXiv: [1512.00325](https://arxiv.org/abs/1512.00325) [astro-ph.HE].
-  Sizun, Patrick, M. Casse, and S. Schanne (2006). “Continuum gamma-ray emission from light dark matter positrons and electrons”. In: *Phys. Rev. D* 74, p. 063514. DOI: [10.1103/PhysRevD.74.063514](https://doi.org/10.1103/PhysRevD.74.063514). arXiv: [astro-ph/0607374](https://arxiv.org/abs/astro-ph/0607374).